

The *NuSTAR* extragalactic surveys: Source catalogs from the Extended *Chandra* Deep Field-South and the *Chandra* Deep Field-North

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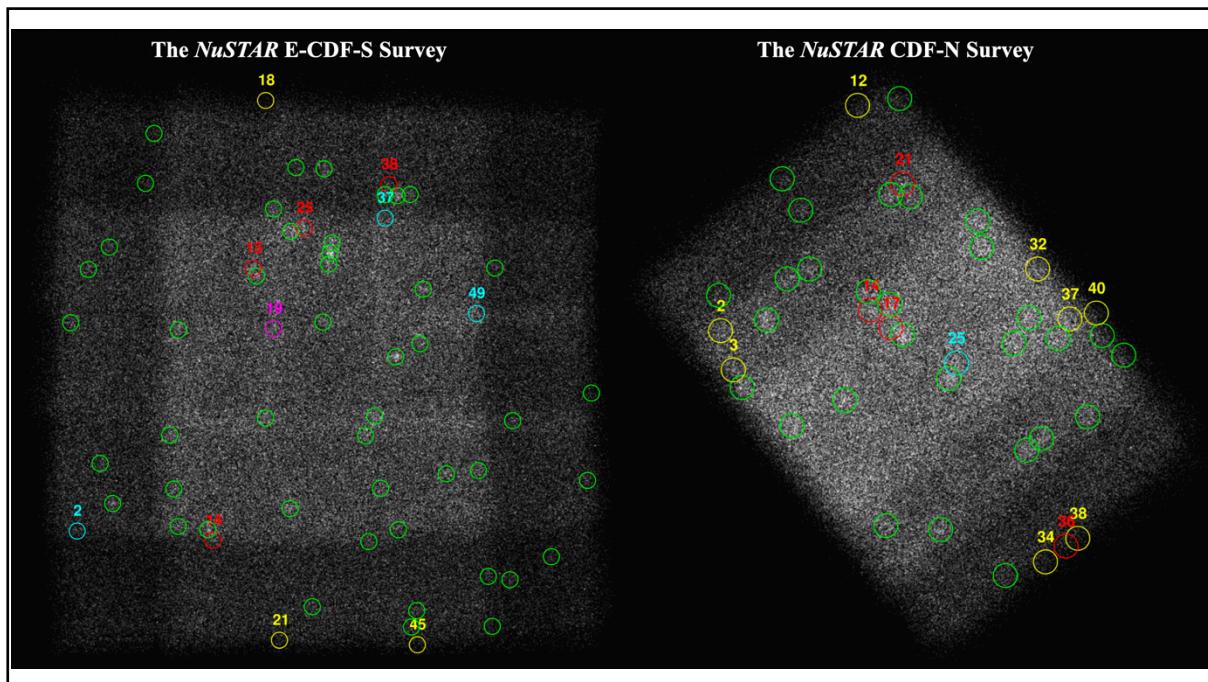
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Graphical abstract



Stacked *NuSTAR* E-CDF-S and CDF-N science mosaics in the 3–24 keV band.

Public summary

- We present a routinized and reliable method to obtain source catalogs from the Nuclear Spectroscopic Telescope Array (*NuSTAR*) extragalactic surveys of the Extended *Chandra* Deep Field-South (E-CDF-S) and *Chandra* Deep Field-North (CDF-N).
- There are 58 and 42 sources in our *NuSTAR* E-CDF-S and CDF-N catalogs, respectively, with the CDF-N catalog being produced for the first time.
- We make our E-CDF-S and CDF-N catalogs publicly available, thereby providing a uniform platform that facilitates further studies involving these two fields.

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Supporting Information

Abstract: We present a routinized and reliable method to obtain source catalogs from the Nuclear Spectroscopic Telescope Array (*NuSTAR*) extragalactic surveys of the Extended *Chandra* Deep Field-South (E-CDF-S) and *Chandra* Deep Field-North (CDF-N). The *NuSTAR* E-CDF-S survey covers a sky area of $\approx 30' \times 30'$ to a maximum depth of ≈ 230 ks corrected for vignetting in the 3–24 keV band, with a total of 58 sources detected in our E-CDF-S catalog; the *NuSTAR* CDF-N survey covers a sky area of $\approx 7' \times 10'$ to a maximum depth of ≈ 440 ks corrected for vignetting in the 3–24 keV band, with a total of 42 sources detected in our CDF-N catalog that is produced for the first time. We verify the reliability of our two catalogs by crossmatching them with the relevant catalogs from the *Chandra* X-ray observatory and find that the fluxes of our *NuSTAR* sources are generally consistent with those of their *Chandra* counterparts. Our two catalogs are produced following the exact same method and made publicly available, thereby providing a uniform platform that facilitates further studies involving these two fields. Our source-detection method provides a systematic approach for source cataloging in other *NuSTAR* extragalactic surveys.

Keywords: Nuclear Spectroscopic Telescope Array; extragalactic survey; X-ray sources; Extended *Chandra* Deep Field-South; *Chandra* Deep Field-North

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1 Introduction

Extragalactic X-ray surveys are efficient at identifying and characterizing highly reliable and fairly complete samples of active galactic nuclei (AGNs), given the following reasons (see, e.g., Brandt and Hasinger^[1], Brandt and Alexander^[2], Xue^[3], for reviews). First, X-ray emission is a nearly universal feature of luminous AGNs, which can be produced in various accretion disk (plus corona) models for AGNs (e.g., Yuan and Narayan^[4]). Second, X-ray emission, especially hard X-ray emission (≥ 10 keV), can penetrate through materials with hydrogen column densities even up to $N_{\text{H}} \approx 10^{25} \text{ cm}^{-2}$, which is key to excavating the majority of the AGN family, i.e., highly obscured and even Compton-thick AGNs (e.g., Li et al.^[5,6]). Third, X-ray emission is subject to minimal dilution by host-galaxy stellar emission and is powerful for probing the immediate vicinity of supermassive black holes (SMBHs) in AGNs even at high redshifts. Last, the production of the X-ray spectrum goes through numerous line and continuum emission processes, and a high-quality X-ray spectrum is effective to infer physical conditions near the central SMBH.

X-ray surveys have resolved a very large portion ($\approx 80\% - 90\%$) of the cosmic X-ray background (CXRB) up

to ≈ 10 keV, with AGNs being the dominant contributor (e.g., Hickox and Markevitch^[7], Xue et al.^[8,9], Lehmer et al.^[10], Luo et al.^[11]), but the resolved fraction around the peak of the CXRB at $\approx 20 - 40$ keV has been very low ($< 10\%$; see, e.g., Brandt and Hasinger^[1], Brandt and Alexander^[2], Harrison et al.^[12]). The Nuclear Spectroscopic Telescope Array (*NuSTAR*), the first focusing high-energy X-ray (3–79 keV) telescope in orbit, has largely broadened the window of X-ray observations^[13]. The *NuSTAR* surveys have resolved $\approx 33\% - 39\%$ of the 8–24 keV CXRB^[12], thereby helping us better understand the contribution of highly obscured and Compton-thick AGNs to the CXRB.

The *Chandra* Deep Fields (CDFs), consisting of the *Chandra* Deep Field-South (CDF-S, Luo et al.^[11], hereafter L17), *Chandra* Deep Field-North (CDF-N, Xue et al.^[14], hereafter X16), and Extended *Chandra* Deep Field-South (E-CDF-S, X16), are important sky areas for study, e.g., AGN demography, physics, and evolution^[3]. The *Chandra* X-ray observatory has accumulated ≈ 7 Ms exposure in the CDF-S (L17), the deepest X-ray exposure ever made, which provides a large sample of AGNs at $z \approx 0 - 5$ for powerful statistical studies. As a parallel field to the CDF-S and being the second deepest X-ray survey, the 2 Ms CDF-N (X16) effectively complements

the 7 Ms CDF-S, accounting for cosmic variance and enabling comparative studies between fields. *NuSTAR* has also observed the CDFs for complementary studies over 10 keV, and has completed a series of additional extragalactic hard X-ray surveys^[12]. Mullaney et al.^[15] (hereafter M15) has released a source catalog from the *NuSTAR* E-CDF-S survey; however, the *NuSTAR* CDF-N catalog is still absent.

In this work, we propose a uniform and reliable method to process the *NuSTAR* E-CDF-S and CDF-N observations and perform source detection. Referring to the previous *NuSTAR* E-CDF-S cataloging work (M15), we obtain both the *NuSTAR* E-CDF-S and CDF-N source catalogs in a routinized and unified way. We describe the production of both catalogs in Section 2 and Section 3, where, for brevity, the data reduction and source detection are introduced in detail only for the E-CDF-S. We summarize our results in Section 4. We use J2000.0 coordinates and a cosmology of $H_0 = 71 \text{ km} \cdot \text{s}^{-1} \cdot \text{Mpc}^{-1}$, $\Omega_M = 0.27$, and $\Omega_\Lambda = 0.73$.

2 Production of the *NuSTAR* E-CDF-S point-source catalog

2.1 Data reduction

We collect 33 valid observations from the *NuSTAR* E-CDF-S survey that cover a sky area of $\approx 30' \times 30'$, almost each of which has an effective exposure of $\approx 45 \text{ ks}$. The details of these observations are presented in Table 1.

2.1.1 Flaring episodes

As *NuSTAR* is composed of two focal plane modules (i.e., FPMA and FPMB), each of the 33 observations results in two event files. We use the program *nupipeline* of the *NuSTAR* data analysis software NuSTARDAS to generate 66 initial event files with default parameters. Following M15, full-field lightcurves in the entire energy band (i.e., 3–78 keV) with a bin size of 20 s are produced to inspect the influence of flaring events. The *dmgti* tool of the *Chandra* interactive analysis of observations (CIAO) is used to make a user-defined good-time interval (GTI) file to avoid background flaring when the average binned count rate exceeds $1.5 \text{ cts} \cdot \text{s}^{-1}$ in the light curves. Taking the GTI files into account, we run *nupipeline* again to obtain the 66 cleaned event files. Following Alexander et al.^[16], the final cleaned event files are split into three standard energy bands, 3–8 keV (soft band; S), 8–24 keV (hard band; H), and 3–24 keV (full band; F), respectively.

2.1.2 Science, exposure, and background mosaics

From the cleaned event files, we produce exposure maps with the NuSTARDAS program *nuexpomap*. For the effects of vignetting, the same energy correction values as those in M15 are adopted to generate the effective exposure maps, i.e., 5.42 keV, 13.02 keV, and 9.88 keV for the soft, hard, and full bands, respectively. The E-CDF-S reaches a maximum depth of $\approx 230 \text{ ks}$ corrected for vignetting in the full band.

Due to the high count-rate backgrounds in the *NuSTAR* E-CDF-S observations, we generate model background maps using the IDL software *nuskybgd*^[17]. Following a similar

strategy adopted by M15, we choose 4 large (i.e., radius of $3'$) circular regions centered on the 4 chips of the detector as our background regions. With the user-defined regions, the *nuskybgd* software can extract and fit the corresponding spectra in XSPEC with the preset models and derive the best-fit parameters. These parameters are used to generate “fake” background images of the observations. Using the FTOOLS task XIMAGE, these simulated images are collected and merged into background mosaics weighted by the corresponding exposure maps; similarly, using XIMAGE, the stacked science mosaics (see Fig. 1) are directly produced from the cleaned event files.

We note that in the newest version of *nuskybgd*, the use of the “nuabs” XSPEC model has been phased out of *nuskybgd* routines. However, we find that when “nuabs” is removed, the model background counts are significantly lower than what they should be. Consequently, we turn to using the old version that includes the “nuabs” model in the spectral fitting process.

2.2 Source detection

As shown in Fig. 1, traditional source detection methods (e.g., WAVDETECT^[18] and ACIS Extract^[19] adopted in Xue et al.^[8, 14], Luo et al.^[11]) are invalid due to the heavy background. Following the general strategy adopted for *NuSTAR* surveys (M15, Masini et al.^[20]), we use the incomplete Gamma (igamma) function (see Georgakakis et al.^[21]) in the *Scipy.special* package to produce false probability (P_{false}) maps for source detection:

$$P_{\text{false}} = \text{igamma}(N_{\text{Sci}}, N_{\text{Bgd}}),$$

where N_{Sci} and N_{Bgd} represent the photon counts within one region at the same position in the science and background mosaics. The P_{false} value gives the probability that a signal with N_{Sci} counts is purely due to random fluctuation given the background of N_{Bgd} , which means that the signal is more likely to be real as P_{false} decreases.

We smooth the science and background mosaics with top-hat functions of different radii, with the former ($10''$) being smaller for finer structures and the latter ($20''$) being larger to decrease the background influence. The P_{false} maps are produced using three methods: (i) the P_{false} value at the position (x, y) is directly derived by $\text{igamma}(\text{Sci}(x, y), \text{Bgd}(x, y))$, where the resulting P_{false} maps are called P_{r0} maps. (ii) At position (x, y) , we perform aperture photometry with a circular region of radius $10''$ on the mosaics and then calculate the P_{false} value from $\text{igamma}(\text{Sci}_{10''}(x, y), \text{Bgd}_{10''}(x, y))$, where the resulting P_{false} maps are called P_{r10} maps. (iii) The same procedure as method (ii) but using a $20''$ radius aperture is adopted to obtain the P_{r20} maps. Considering the potential signals residing in the local minima of the P_{false} maps, we produce the inverse P_{false} maps using $\log(1/P_{\text{false}})$ to identify peaks with the SExtractor source-detection algorithm^[22].

Some modifications on the default SExtractor configuration file *default.sex* are listed below:

(I) DETECT_MINAREA is set to 1. Because we smooth the input maps, even one pixel in the $1/P_{\text{false}}$ maps can be considered as a potential source.

Table 1. Details of the *NuSTAR* E-CDF-S observations.

Obs. ID	Obs. Name	Obs. Date	RA	DEC	t_{eff}
60022001002	ECDFS_MOS001	2012-09-28	52.93	-27.97	49.0
60022002001	ECDFS_MOS002	2012-09-29	52.93	-27.97	50.3
60022003001	ECDFS_MOS003	2012-09-30	52.93	-27.97	50.2
60022004001	ECDFS_MOS004	2012-10-01	52.93	-27.97	50.9
60022005001	ECDFS_MOS005	2012-10-02	53.06	-27.86	50.5
60022006001	ECDFS_MOS006	2012-10-04	53.06	-27.86	49.2
60022007002	ECDFS_MOS007	2012-11-30	53.06	-27.86	51.7
60022008001	ECDFS_MOS008	2012-12-01	53.06	-27.86	51.7
60022009001	ECDFS_MOS009	2012-12-03	53.18	-27.75	50.3
60022010001	ECDFS_MOS010	2012-12-04	53.18	-27.75	51.2
60022011001	ECDFS_MOS011	2012-12-05	53.18	-27.75	51.7
60022012001	ECDFS_MOS012	2012-12-06	53.18	-27.75	52.1
60022013001	ECDFS_MOS013	2012-12-07	53.31	-27.64	52.5
60022014001	ECDFS_MOS014	2012-12-08	53.31	-27.64	52.9
60022015001	ECDFS_MOS015	2012-12-09	53.31	-27.64	53.2
60022016001	ECDFS_MOS016	2012-12-10	53.31	-27.64	50.1
60022016003	ECDFS_MOS016	2013-03-15	52.93	-27.64	51.7
60022015003	ECDFS_MOS015	2013-03-17	52.93	-27.64	51.2
60022014002	ECDFS_MOS014	2013-03-18	52.93	-27.64	51.4
60022013002	ECDFS_MOS013	2013-03-19	52.93	-27.64	49.7
60022012002	ECDFS_MOS012	2013-03-20	53.06	-27.75	49.7
60022011002	ECDFS_MOS011	2013-03-21	53.06	-27.75	48.9
60022010002	ECDFS_MOS010	2013-03-22	53.06	-27.75	32.5
60022010004	ECDFS_MOS010	2013-03-23	53.06	-27.75	16.4
60022009003	ECDFS_MOS009	2013-03-24	53.18	-27.75	49.5
60022008002	ECDFS_MOS008	2013-03-25	53.18	-27.86	49.8
60022007003	ECDFS_MOS007	2013-03-26	53.18	-27.86	49.6
60022006002	ECDFS_MOS006	2013-03-27	53.18	-27.86	48.7
60022005002	ECDFS_MOS005	2013-03-28	53.31	-27.86	49.3
60022004002	ECDFS_MOS004	2013-03-29	53.31	-27.97	48.8
60022003002	ECDFS_MOS003	2013-03-30	53.31	-27.97	49.0
60022002002	ECDFS_MOS002	2013-03-31	53.31	-27.97	49.0
60022001003	ECDFS_MOS001	2013-04-01	52.93	-27.97	48.2

Obs. ID is a unique identification number specifying the *NuSTAR* observation; Obs. Name gives the designation of the target at which *NuSTAR* was pointing. Obs. Date is the start time of the observation. RA and DEC give the J2000.0 Right Ascension and the Declination of the *NuSTAR* pointing position. t_{eff} is the effective exposure time (in ks) after background filtering (see Section 2.1.1).

(II) THRESH_TYPE is set to ABSOLUTE. Under this condition, DETECT_THRESH represents the detection threshold above which a signal in the maps can be considered significant.

(III) FILTER and CLEAN are set to N(o). We do not need these processes as what we address here are not real observational images.

With proper DETECT_THRESH values, SExtractor is able to detect potential sources in the three standard bands to produce our initial catalogs.

2.2.1 Simulations

We perform a series of simulations to determine DETECT_THRESH. Because the science mosaics are smoothed with the 10''-radius circular top-hat function (see Section 2.2), we split the background images into several square cells with a side length of 20''. For every background image, Poisson realizations are performed on these cells to make up a “science” image from pure Poisson fluctuation. The aforementioned procedures are then performed on these Poisson realizations, and SExtractor should detect no sources in the corresponding

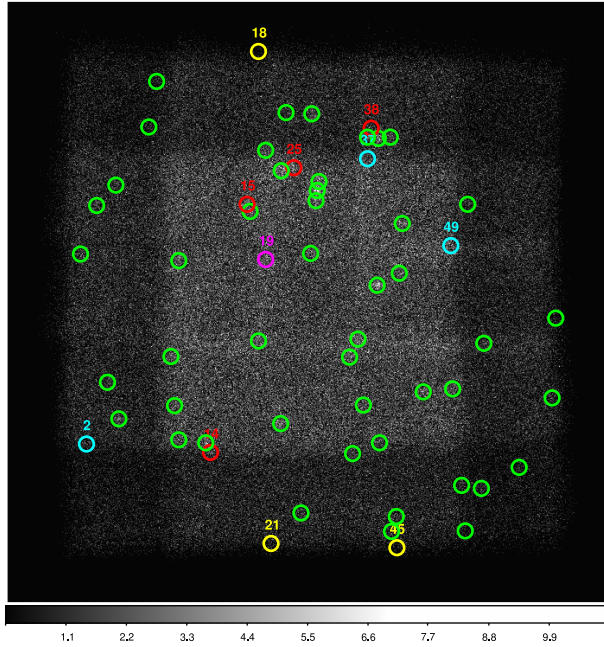


Fig. 1. Stacked *NuSTAR* E-CDF-S science mosaic in the full band, with a total of 58 sources plotted as circles. The green (47/58) and red (4/58; being less significant detections) sources have the *Chandra* 250 ks E-CDF-S counterparts within $r_m = 30''$. The magenta source (1/58) does not have any *Chandra* 250 ks E-CDF-S counterparts but can be matched to a *Chandra* 7 Ms CDF-S source. Among the unmatched sources (6/58), 3 yellow sources reside in the very edges of the mosaic and 3 cyan sources reside in the chip-gap areas. Numbers are source XIDs. The bottom color bar indicates the counts per pixel.

inverse P_{false} maps with proper thresholds.

The simulations are repeated for 100 times, in each of which we produce inverse P_{false} maps with three different aperture radii (i.e., $0''$, $10''$, and $20''$) in the three standard bands. DETECT_THRESH is set to the value above which SExtractor can find no more than N signals in these simulated maps per aperture radius per band, and $R = N/100$ represents the false detection rate. These thresholds are applied to the real inverse P_{false} maps for source detection.

2.2.2 Final catalog production

We first set $R = 0.1$ for loose thresholds, with which our algorithm is run on the P_{r0} , P_{r10} , and P_{r20} maps to generate three seed lists in each band. These seed lists are produced from different P_{false} maps; thus, there might be multiple nearby detections belonging to the same source. To identify every unique source, we merge the seed lists for each band and then

run a modified friends-of-friends algorithm (hereafter FOF) on these merged lists for deduplication.

FOF is common in cosmology for identifying groups in dense fields^[23], which demands that any particle (or galaxy) that finds another one within a distance l (called linking length) should be linked to it to form a group. Sources in the merged lists are split into group sources and isolated sources by FOF with a linking length of $30''$. For each group, members are ranked by their inverse P_{r20} values (even if being initially detected from other P_{false} maps), then compete against any others within a $30''$ radius, in which a member of higher value can survive and participate in a next competition. After several rounds of competitions, the final survivals from the group sources are combined with the isolated sources to make up the deduplicated seed lists for each band.

The deduplicated seed lists from different bands are combined into one seed list and then split by FOF in the same way. For each group in the combined list, members are ranked by their inverse P_{r20} values in the full band (regardless of whether being detected in this band), then participate in competitions against each other. After deduplication, the remaining sources are collected to construct our seed catalog with $R = 0.1$. The details are presented in Table 2.

To produce a reliable final catalog, we then set $R = 0.01$, which means only 1 false signal being detected in 100 simulations. The thresholds of the inverse P_{r20} maps are 3.99, 4.36, and 4.33 in the soft, hard, and full bands, respectively, corresponding to $\approx 99\%$ reliability^[24]. We remove the sources that do not meet any of our final thresholds and construct the final catalog with the remaining ones.

The E-CDF-S final catalog contains 58 sources, each of which is detected in at least one of the three standard bands. Of these 58 sources, 33, 13, and 54 are detected in the soft, hard, and full bands, respectively; 3, 1, and 21 are detected only in the soft, hard, and full bands, respectively; no source is detected in exactly the soft and hard bands, 21 in exactly the soft and full bands, and 3 in exactly the hard and full bands; and 9 are detected in all the three standard bands.

2.2.3 Photometry and deblending

The radius of the 90% encircled-energy fraction contour of the *NuSTAR* point spread function (PSF) is approximately $67.5''$, which is relatively large compared to the average distance among sources. We adopt a similar strategy to that of M15 to choose an aperture size of $30''$ for photometry extraction and assume that the net counts within this aperture are only contaminated by other nearby *NuSTAR*-detected sources within $90''$.

Table 2. The *NuSTAR* E-CDF-S cataloging process.

Seed catalog ($R = 0.1$)	Source # (P_{r0})	Source # (P_{r10})	Source # (P_{r20})	Source # (Total)	Source # (After deduplication)
3–8 keV	33	35	35	103	46
8–24 keV	17	13	20	50	23
3–24 keV	42	41	55	138	66
Final catalog ($R = 0.01$ & P_{r20})	Source # (3–8 keV)	Source # (8–24 keV)	Source # (3–24 keV)	Source # (3–8 keV & 8–24 keV)	Source # (Total)
Before deblending	33	13	54	9	58
After deblending	33	13	50	9	54

For each source in our catalog, the total and background counts are calculated within a circular region of radius $30''$ in the science and background mosaics, respectively, while the net counts are derived by subtracting the background counts from the total counts. Following Gehrels^[25], we estimate the upper and lower 1σ confidence limits on the total counts; for those not detected in certain bands, only the upper limits are derived, using:

$$\sum_{x=0}^n \frac{\lambda_u^x e^{-\lambda_u}}{x!} = 1 - \text{CL},$$

$$\sum_{x=0}^{n-1} \frac{\lambda_l^x e^{-\lambda_l}}{x!} = \text{CL},$$
(1)

where λ_u and λ_l represent the upper and lower limits, n is the photon count, and CL represents the confidence level, respectively.

The background count error can be approximated by $\sigma_{\text{Bgd}} = \frac{1 + \sqrt{\text{factor} \times C_{\text{Bgd}} + 3/4}}{\text{factor}}$, where C_{Bgd} is the background count and the factor gives the ratio between the total area where the background model is defined and the area for photometry extraction (i.e., $\text{factor} = (180''/30'')^2 \times 4 = 144$). Subsequently, the upper and lower limits on the net counts are calculated as $\sigma_{\text{Net,u}} = \sqrt{\lambda_u^2 + \sigma_{\text{Bgd}}^2}$ and $\sigma_{\text{Net,l}} = \sqrt{\lambda_l^2 + \sigma_{\text{Bgd}}^2}$, respectively.

To deblend the sources in our catalog, the FOF algorithm is applied again to split them into group sources and isolated sources with a different linking length of $90''$. For the isolated sources, we assume that they cannot be contaminated by any

other sources (away beyond $90''$); for the group sources, a system of n linear simultaneous equations is established:

$$\begin{aligned} C_T^1 &= N(r_{1,1})C_D^1 + N(r_{1,2})C_D^2 + \dots + N(r_{1,n})C_D^n, \\ C_T^2 &= N(r_{2,1})C_D^1 + N(r_{2,2})C_D^2 + \dots + N(r_{2,n})C_D^n, \\ &\vdots \\ C_T^n &= N(r_{n,1})C_D^1 + N(r_{n,2})C_D^2 + \dots + N(r_{n,n})C_D^n, \end{aligned}$$
(2)

where C_T^n is the total net counts of source n , C_D^n is the deblended net counts of source n , and $N(r_{i,j})$ is the normalized function of the separation between the sources i and j ($r_{i,j}$ represents the separation distance, while $N(0) = 1$), in which several simplifications are proposed to avoid the complications of the nonazimuthally symmetric *NuSTAR* PSF.

Following the deblending procedure above, we then perform deblending with another aperture of $20''$ radius, and recalculate the P_{false} of each source after deblending. The postdeblending P_{false} values are compared with the P_{r20} thresholds, and 4 of the 58 sources in our catalog become no longer significant. Additionally, we find 1 source in the area of relatively low exposure (< 40 ks, corresponding to $\lesssim 10\%$ of the maximum survey exposure). All of these 5 sources are detected in the full band only, and we flag but do not remove them (see Fig. 1).

To validate the reliability of our catalog (a total of 58 sources), we match it to the previous *NuSTAR* E-CDF-S catalog (a total of 54 sources) in M15 using a matching radius $r_m = 30''$ and find a total of 36 counterpart pairs. We compare their net counts in Fig. 2 and find good consistency within 1σ errors. We also compare their aperture-corrected fluxes (see Section 2.3) and find good agreement between each

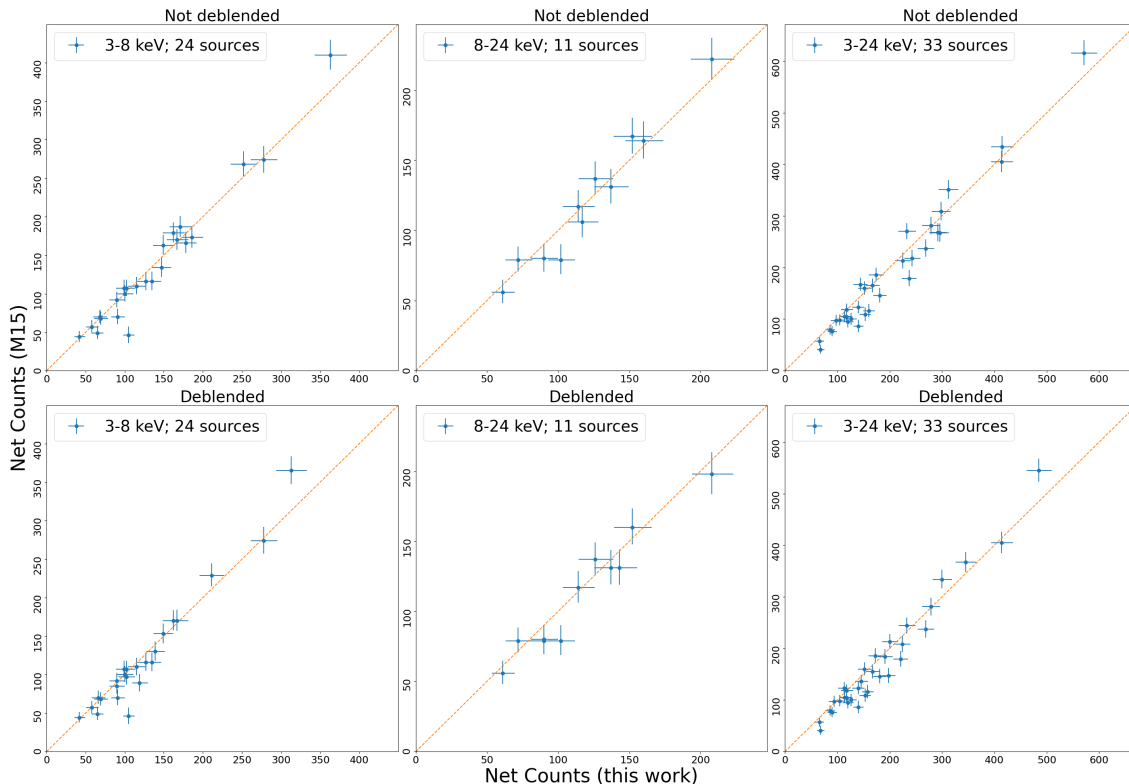


Fig. 2. Comparison of net counts obtained by this work and M15 in the E-CDF-S. The diagonal dashed lines indicate the 1 : 1 relations.

other.

A significant fraction of the M15 sources are not detected by our work (and vice versa), mainly due to two facts: the detailed cataloging methodologies are different between our work and M15, and those unmatched sources generally have lower net counts such that they could be too faint to be detected by either work.

2.3 Matching to the *Chandra* E-CDF-S and CDF-S catalogs

We first match our catalog to the *Chandra* 250 ks E-CDF-S catalog (X16) using $r_m = 30''$, and find 51 of the 58 sources to have at least one *Chandra* counterpart. In these matches, 20, 24, 5, and 2 *NuSTAR* sources have 1, 2, 3, and 4 *Chandra* counterparts, respectively; no *NuSTAR* sources has more than 4 *Chandra* counterparts. For the flagged sources (see Section 2.2.3), all of the 4 less-significant sources (XIDs = 14, 15, 25, 38; see Fig. 1) have *Chandra* counterparts, but the source with low exposure time (XID = 18) does not. We then match the 7 sources without *Chandra* 250 ks counterparts to the *Chandra* 7 Ms CDF-S catalog (L17) using $r_m = 30''$, only finding one further match (XID = 19).

We inspect the positions and properties of the 6 unmatched sources, finding that 3 reside in the very edge of the *NuSTAR* E-CDF-S survey mosaic (XIDs = 18, 21, and 45; see Fig. 1), which implies uncertainty in their detections. For the 3 remaining sources, one (XID = 2) is also near the edge of the mosaic and detected only in the 3–24 keV band, and the other two sources (XIDs = 37, 49) reside in the chip-gap areas (see Fig. 1) which might be spurious. Because we aim to find as many sources as possible using our algorithm alone (without manual intervention), these 6 sources are conserved and flagged in our final catalog.

To compare the fluxes of these matched sources, the observed deblended fluxes of the *NuSTAR* sources are derived

following the same approach as Alexander et al.^[16] For the sources detected in both the soft and hard bands, we calculate their hardness ratios, $HR=(H-S)/(H+S)$, using the Bayesian estimation of hardness ratios method^[26]; for other sources, an HR value corresponding to the power-law spectral photon index of $\Gamma = 1.8$ is assumed. Using the derived HRs, the same parameters as in M15 are then adopted to convert count rates to observed fluxes (see Section 2.3.3 of M15), which reach a soft-band flux limit of $\approx 10^{-14} \text{ erg} \cdot \text{s}^{-1} \cdot \text{cm}^{-2}$.

Due to the different observable energy range of *Chandra* (0.5–7 keV), we assume the *Chandra* counterparts having a simple power-law spectrum (i.e., $f(E) \propto E^{-\Gamma}$), then derive the conversion factor between the fluxes in the 2–7 keV and 3–8 keV bands. The fluxes in the 2–7 keV band and photon indices (Γ) can be obtained from the *Chandra* catalogs (X16; L17). For the multiple *Chandra* counterparts within $r_m = 30''$ of the *NuSTAR* sources, we calculate their total 3–8 keV fluxes instead. The comparison between the *NuSTAR* and total *Chandra* 3–8 keV fluxes of the matched sources is shown in Fig. 3, which indicates general agreement within a factor of 3 for the majority of the sources. However, the *NuSTAR* fluxes appear to be systematically lower than the *Chandra* fluxes, which is mainly due to that the *NuSTAR* measured/assumed photon index may be different from the *Chandra* measured/assumed photon index (in the case of 1-to-1 match) or photon indexes (in the case of 1-to-multiple match).

3 Production of the *NuSTAR* CDF-N point-source catalog

3.1 Data reduction and source detection

We collect 12 valid observations from the *NuSTAR* CDF-N survey that cover a sky area of $\approx 7' \times 10'$, almost each of which has an effective exposure of ≈ 50 ks. We summarize

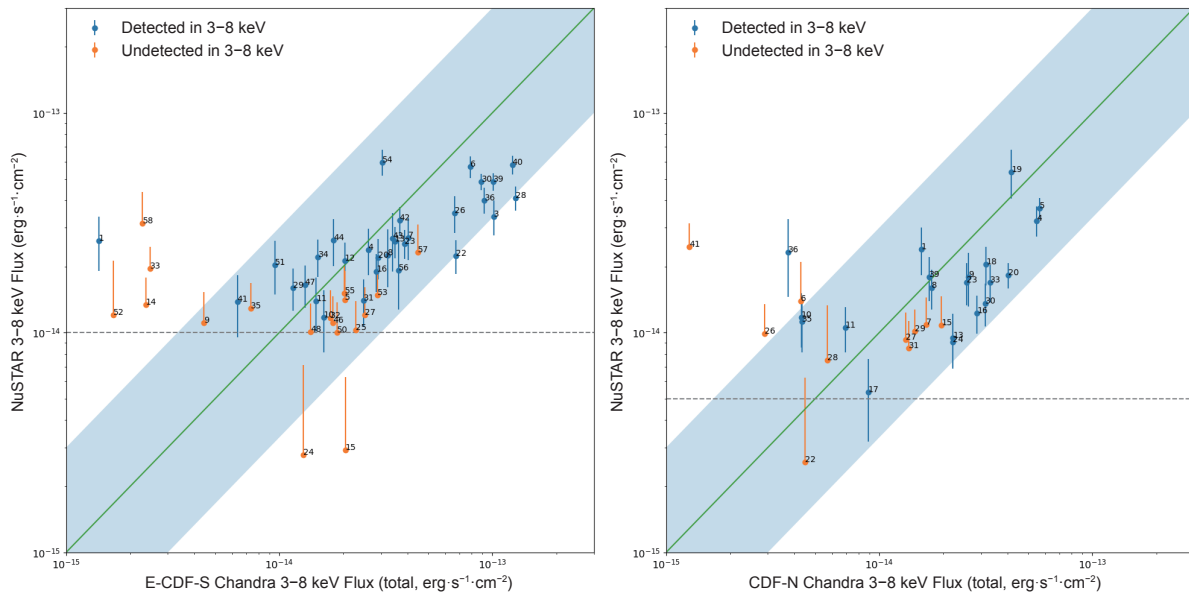


Fig. 3. Comparison between the *NuSTAR* E-CDF-S deblended fluxes (this work) and the total *Chandra* E-CDF-S (X16; left panel) and CDF-S (L17; right panel) fluxes in the 3–8 keV band (all fluxes are aperture-corrected). In each panel, the green 1 : 1 line is centered at the shaded area indicating a factor of ≤ 3 difference from the 1 : 1 line, and the horizontal dotted line indicates the detection limit.

Table 3. Details of the *NuSTAR* CDF-N observations.

Obs. ID	Obs. Name	Obs. Date	RA	DEC	t_{eff}
60110003002	GOODSN_MOS003	2015-04-22	12.61	+62.20	45.6
60110002001	GOODSN_MOS002	2015-04-23	12.62	+62.24	45.6
60110001001	GOODSN_MOS001	2015-04-24	12.62	+62.28	46.1
60110001003	GOODSN_MOS001	2015-08-04	12.62	+62.28	59.7
60110002003	GOODSN_MOS002	2015-08-09	12.61	+62.24	64.3
60110003003	GOODSN_MOS003	2015-08-10	12.61	+62.20	59.0
60110001005	GOODSN_MOS001	2015-10-31	12.62	+62.27	61.5
60110002004	GOODSN_MOS002	2015-11-01	12.62	+62.23	64.3
60110003004	GOODSN_MOS003	2015-11-02	12.61	+62.20	67.8
60110001007	GOODSN_MOS001	2016-01-30	12.62	+62.27	52.8
60110002005	GOODSN_MOS002	2016-01-31	12.62	+62.23	52.9
60110003005	GOODSN_MOS003	2016-02-01	12.61	+62.19	54.0

Parameters' interpretation is the same as those of Table 1.

the information of these observations in Table 3. For data reduction, we apply the exactly same procedures as in Section 2.1, therefore the technical details are introduced briefly in this section.

Full-field lightcurves in the entire energy band with a bin size of 20 s are produced to inspect the influence of flaring events, and a threshold of $1.3 \text{ cts} \cdot \text{s}^{-1}$ is selected to reject the periods with high background flares. Following the same procedures in Section 2.1.2, the science, exposure, and background mosaics are created from the cleaned event files. The CDF-N reaches a maximum depth of $\approx 440 \text{ ks}$ corrected for vignetting in the full band, almost doubling that of the E-CDF-S. We present the stacked science mosaic in the full band in

Fig. 4.

Again we set $R = 0.1$ to obtain the seed list from the P_{false} maps, then refilter these seed sources with a strict threshold of $R = 0.01$ for the final catalog that contains 42 CDF-N sources, each of which is detected in at least one of the three standard bands. Of the 42 sources, 26, 11, and 31 are detected in the soft, hard, and full bands, respectively; 9, 2, and 12 are detected in the soft, hard, and full bands only, respectively; no source is detected in exactly both the soft and hard bands, 10 in exactly both the soft and full bands, and 2 in exactly both the hard and full bands; 7 are detected in all the three standard bands. After deblending, 6 of the 42 sources are no longer significant, and we still flag and keep them in our final catalog. These results are listed in Table 4.

3.2 Matching to the *Chandra* CDF-N catalog

We match our catalog to the *Chandra* 2 Ms CDF-N catalog (X16) using $r_m = 30''$, and find 33 of the 42 sources to have at least one *Chandra* counterpart. In these matches, 14, 11, 4, and 3 *NuSTAR* sources have 1, 2, 3, and 4 *Chandra* counterparts, respectively, and 1 *NuSTAR* source has more than 4 *Chandra* counterparts.

For the *NuSTAR* sources without *Chandra* counterparts, we inspect their positions and properties and find that almost all of them (8/9; XIDs = 2, 3, 12, 32, 34, 37, 38, 40) reside in or near the edges of the *NuSTAR* CDF-N field and the remaining one (XID = 25) is only detected in the hard band which might be too “hard” to be detected by *Chandra* (see Fig. 4).

We compare the fluxes of these matched sources in Fig. 5, which also indicates general agreement within a factor of 3 for the majority of the sources. The normalized flux histograms of the *NuSTAR* CDF-N and E-CDF-S sources are compared in Fig. 6: CDF-N sources generally have lower fluxes than E-CDF-S sources, being consistent with the fact that the average *NuSTAR* CDF-N exposure depth (reaching a soft-band flux limit of $\approx 5 \times 10^{-15} \text{ erg} \cdot \text{s}^{-1} \cdot \text{cm}^{-2}$) is almost twice that of the E-CDF-S.

4 Conclusions

In this work, we collect the original observations from the

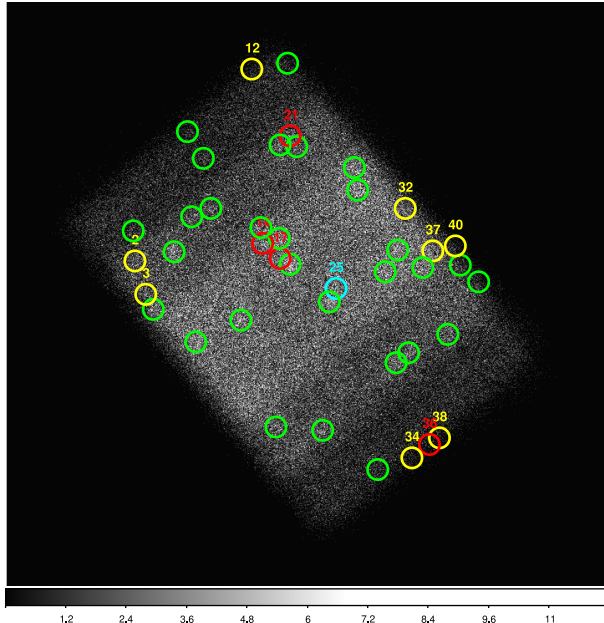
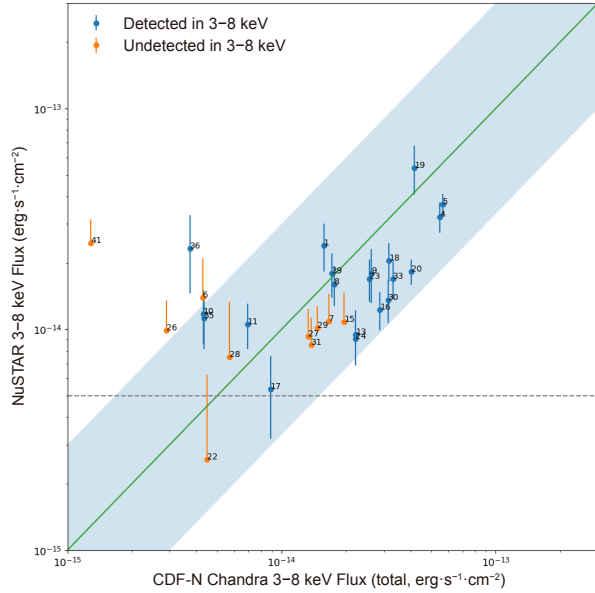


Fig. 4. Stacked *NuSTAR* CDF-N science mosaic in the 3–24 keV band (cf. Fig. 1), with a total of 42 sources plotted. The green (29/42) and red (4/42; being less significant detections) sources have the *Chandra* 2 Ms CDF-N counterparts within $r_m = 30''$. Among the unmatched sources (9/42), 8 yellow sources reside in or near the edges of the mosaic, and 1 cyan source is only detected in the 8–24 keV band.

Table 4. The *NuSTAR* CDF-N cataloging process.

Seed catalog ($R = 0.1$)	Source # (P_{r0})	Source # (P_{r10})	Source # (P_{r20})	Source # (Total)	Source # (After deduplication)
3–8 keV	27	19	28	74	32
8–24 keV	4	9	13	26	14
3–24 keV	42	26	30	98	38
Final catalog ($R = 0.01$ & P_{r20})	Source # (3–8 keV)	Source # (8–24 keV)	Source # (3–24 keV)	Source # (3–8 keV & 8–24 keV)	Source # (Total)
Before deblending	26	11	31	7	42
After deblending	22	11	27	6	36

**Fig. 5.** Same as Fig. 3, but for the comparison between the *NuSTAR* CDF-N deblended fluxes (this work) and the total *Chandra* CDF-N fluxes (X16) in the 3–8 keV band (all fluxes are aperture-corrected).

NuSTAR E-CDF-S and CDF-N surveys and produce cleaned event files. Simulated background mosaics are generated

using the IDL software *nuskybgd*, and then processed along with science mosaics to produce P_{false} maps for source detection.

For the *NuSTAR* E-CDF-S survey, the main results are as follows:

(I) The E-CDF-S catalog consists of 58 sources that are detected using our algorithm without manual intervention, with 4 of them being not significant after deblending.

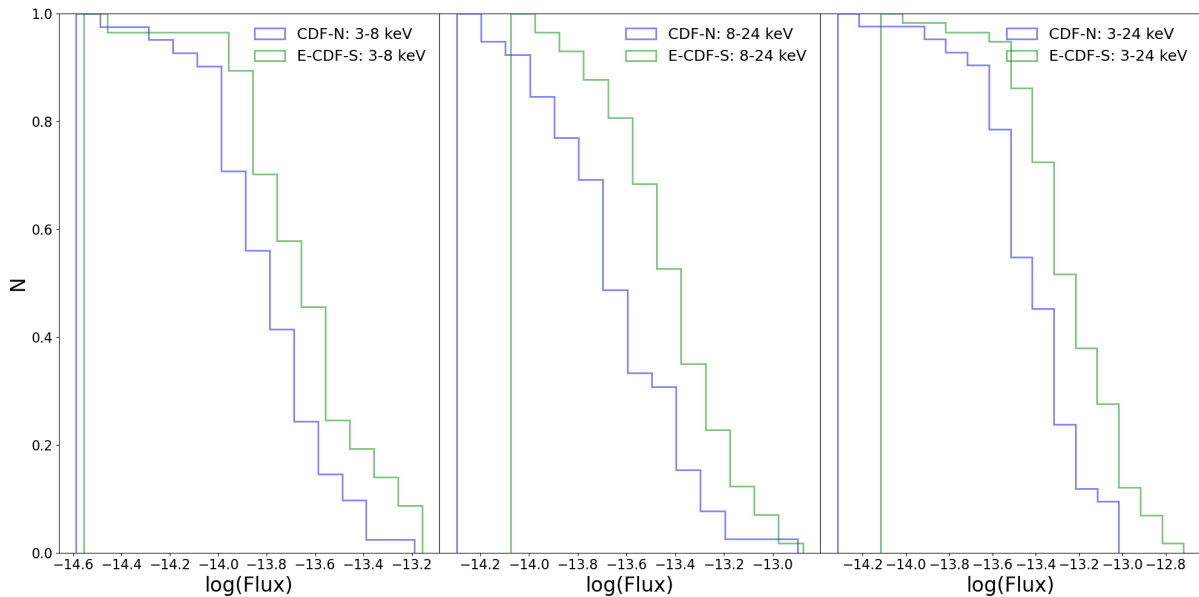
(II) We compare our catalog with the previous *NuSTAR* E-CDF-S catalog (M15) using $r_m = 30''$ and find a total of 36 matches, the net counts of which agree well within 1σ errors.

(III) We compare our catalog with the *Chandra* E-CDF-S and CDF-S catalogs (X16 and L17), and find a total of 51 matches, the fluxes of which agree well above the detection limit. All of the 4 sources being not significant after deblending have counterparts in the *Chandra* catalogs; and the 7 unmatched sources are flagged as being spurious but still conserved in the catalog.

For the *NuSTAR* CDF-N survey, the main results are as follows:

(I) The CDF-N catalog, produced for the first time by this work, consists of 42 sources that are detected using our algorithm without manual intervention, with 6 of them being not significant after deblending.

(II) We compare our catalog with the *Chandra* CDF-N

**Fig. 6.** Normalized distributions of deblended *NuSTAR* fluxes in the CDF-N and E-CDF-S.

catalog (X16), and find a total of 33 matches, the fluxes of which agree well above the detection limit. We flag the 9 unmatched sources as being spurious but conserve them in the catalog.

(III) The flux limits are significantly lower in the *NuSTAR* CDF-N field (having deeper exposures) than that in the *NuSTAR* E-CDF-S field.

Our source-detection method provides a systematic approach for source cataloging in other *NuSTAR* extragalactic surveys. We make our *NuSTAR* E-CDF-S and CDF-N source catalogs publicly available (see Appendix for catalog description), which provide a uniform platform that facilitates further studies involving these two fields.

Supporting information

The supporting information for this article can be found online at <https://doi.org/10.52396/JUSTC-2023-0032>. The supporting information includes the CDF-N and E-CDF-S fits.

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Conflict of interest

The authors declare that they have no conflict of interest.

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Appendix

The *NuSTAR* E-CDF-S and CDF-N source catalogs have the same 51 columns that are summarized in Table A1, with the details described below.

- (1) Column 1 gives the source sequence number (i.e., XID). We list the sources in the order of decreasing right ascension.
- (2) Columns 2 and 3 give the right ascension and declination of the *NuSTAR* source, respectively.
- (3) Column 4 gives the flag whether the exposure time of the source is above the threshold of ≈ 40 ks.
- (4) Columns 5–7 give the flags indicating in which of the three standard bands (3–8 keV, 8–24 keV, and 3–24 keV bands) the source is detected.
- (5) Columns 8–10 give the flags indicating in which of the three standard bands the source is detected after deblending.
- (6) Columns 11–13 give the P_{false} values for the three standard bands.
- (7) Columns 14–16 give the P_{false} values for the three standard bands after deblending.
- (8) Column 17 gives the flag whether the source remains significant after deblending.
- (9) Columns 18–20 give the post-deblending net counts with the corresponding lower and upper errors in the 3–8 keV band.
- (10) Columns 21–23 give the post-deblending net counts with the corresponding lower and upper errors in the 8–24 keV band.
- (11) Columns 24–26 give the post-deblending net counts with the corresponding lower and upper errors in the 3–24 keV band.
- (12) Columns 27–29 give the effective exposure times derived from the exposure maps for the three standard bands.
- (13) Columns 30–32 give the post-deblending flux with the corresponding lower and upper errors in units of $10^{-14} \text{ erg} \cdot \text{s}^{-1} \cdot \text{cm}^{-2}$ in the 3–8 keV band (hereafter the same units).
- (14) Columns 33–35 give the post-deblending flux with the corresponding lower and upper errors in the 8–24 keV band.
- (15) Columns 36–38 give the post-deblending flux with the corresponding lower and upper errors in the 3–24 keV band.
- (16) Columns 39–41 give the hardness ratio with the corresponding lower and upper errors.
- (17) Columns 42–46 give the *Chandra* counterpart with the corresponding catalog name, source sequence number, right ascension and declination, and distance (in units of arcsecond) from the *NuSTAR* source.
- (18) Column 47 gives the flux of the *Chandra* counterpart in the 3–8 keV band.
- (19) Column 48 gives the total flux of the *Chandra* counterparts found within $30''$ in the 3–8 keV band.
- (20) Columns 49 and 50 give the photometric redshift and spectroscopic redshift of the *Chandra* counterpart.
- (21) Column 51 gives the adopted redshift of the source, with the spectroscopic redshift preferred if available.

Table A1. Overview of columns in the *NuSTAR* E-CDF-S and CDF-N catalogs.

Column(s)	Description
1	Source sequence number (i.e., XID) in this work
2, 3	J2000.0 RA and DEC of the <i>NuSTAR</i> source
4	Flag of whether the exposure time of the source is above the threshold
5–7	Flags indicating in which of the three standard bands the source is detected
8–10	Flags indicating in which of the three standard bands the source is detected after deblending
11–13	P_{false} values for the three standard bands
14–16	P_{false} values for the three standard bands after deblending
17	Flag of whether the source remains significant after deblending
18–26	Net counts and associated errors in the three standard bands after deblending
27–29	Effective exposure times in the three standard bands
30–38	Fluxes and associated errors in the three standard bands after deblending
39–41	Hardness ratio and associated errors
42–46	Catalog name, sequence number, RA, and DEC of the <i>Chandra</i> counterpart, separation between the <i>NuSTAR</i> position and the <i>Chandra</i> counterpart
47	Flux of the <i>Chandra</i> counterpart in the 3–8 keV band
48	Total fluxes of the <i>Chandra</i> counterparts within $30''$ in the 3–8 keV band
49, 50	Photometric redshift and spectroscopic redshift of the <i>Chandra</i> counterpart
51	Adopted redshift of the <i>NuSTAR</i> source